Gravitational-Wave Transient Astronomy: the next 20 years

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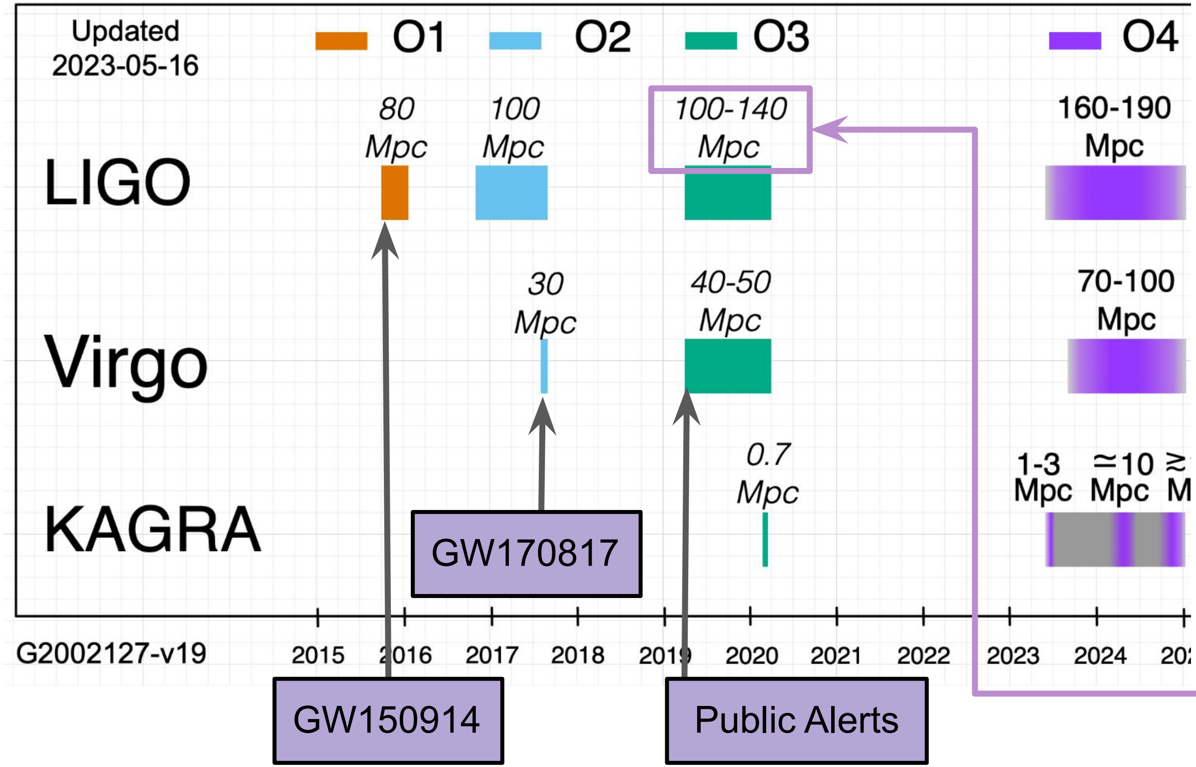
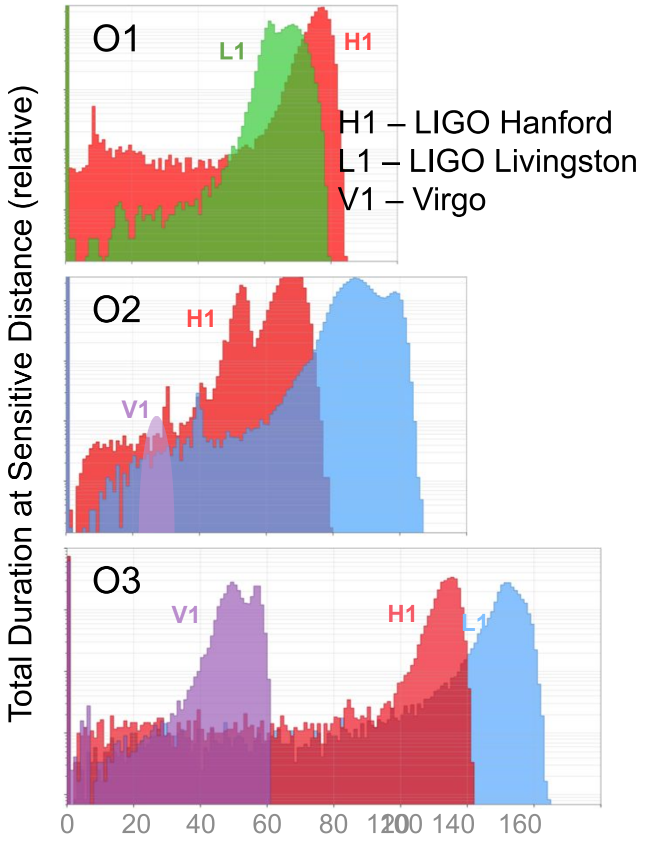
The Transient and Variable Universe
20 June 2023

https://dcc.ligo.org/G2301194

International Gravitational-Wave Observatory Network (IGWN)

Observing runs

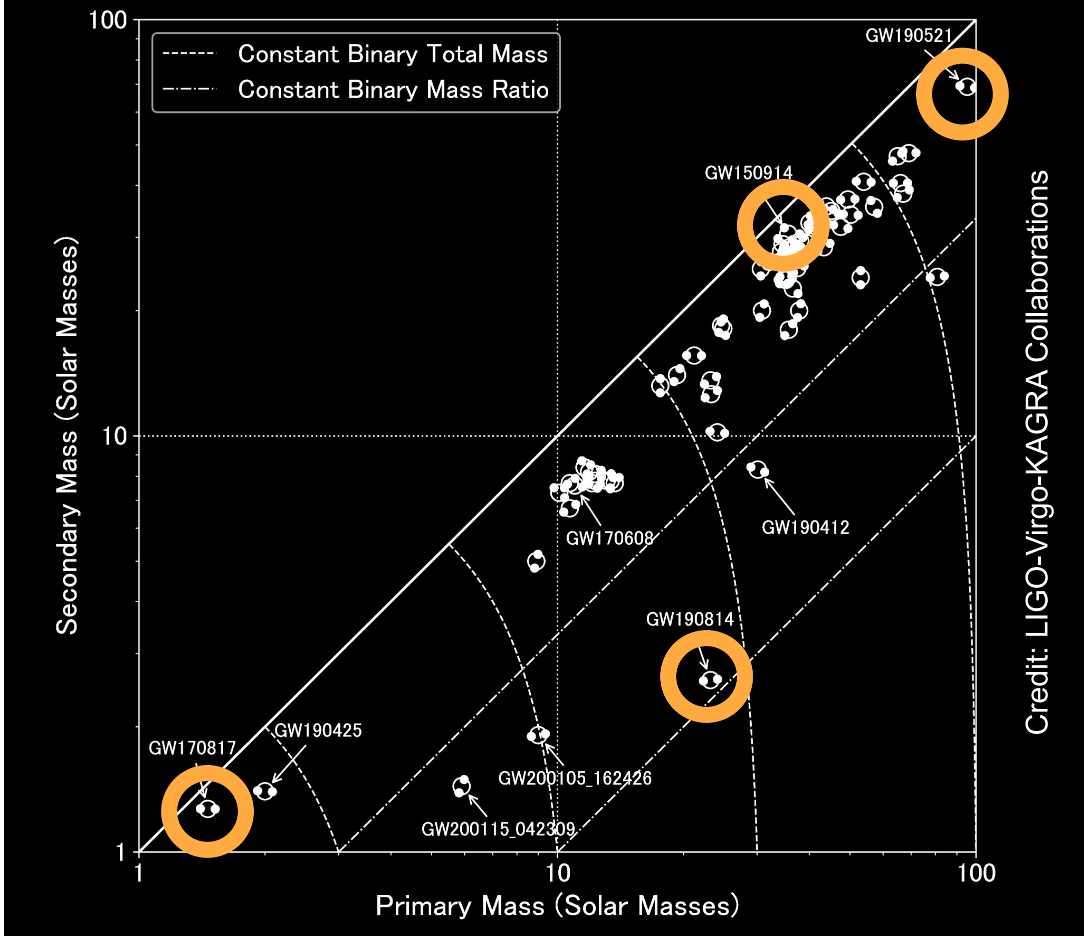




Angle Averaged Binary Neutron Star Inspiral Range

Mergers

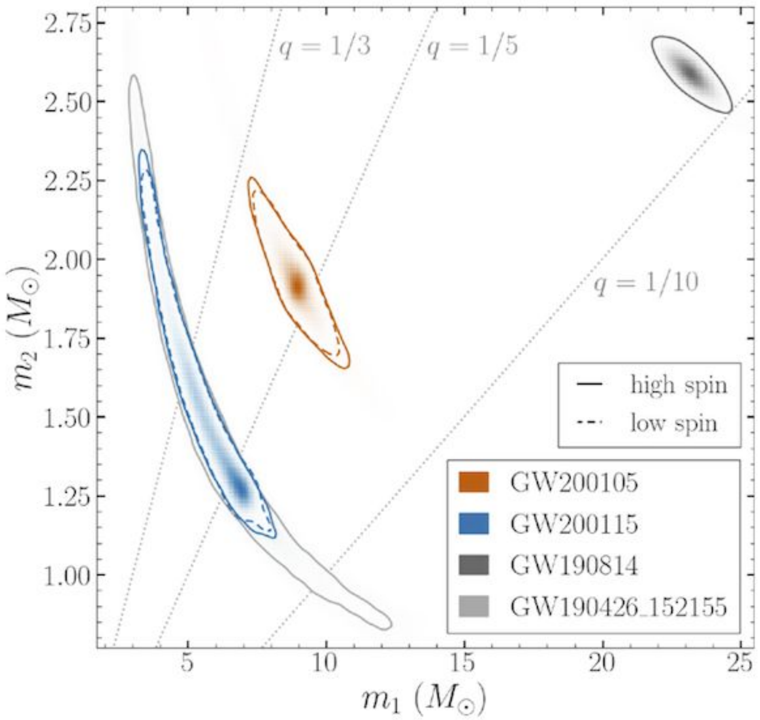
- GW150914
  - First astrophysical source
  - Binary black holes exist
- GW170817
  - Binary neutron star mergers are gamma-ray burst progenitors
- GW190521
  - Black holes exist in pair instability mass gap
- GW190814
  - Compact objects exist with masses between 2-5 Msun



Mergers involving neutron stars

- GW170817 & GW190425
  - Binary neutron star (BNS) merger waves
- GW170817 & GRB 170817A
  - Fractional difference in speed of gravity and the speed of light is between -3 x 10^{-15} and 7 x 10^{-16}
- GW170817 & AT 2017gfo
  - Binary neutron star mergers produce kilonova explosions that generate heavy elements

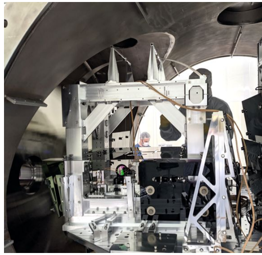
B. P. Abbott et al 2017 ApJL 848 L13

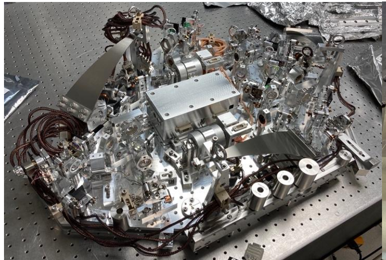


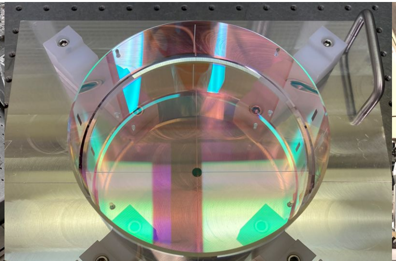
R. Abbott *et al* 2021 *ApJL* **915** L5

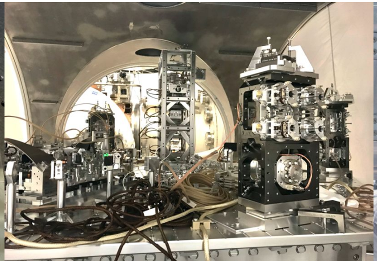
LIGO instrumental upgrades for O4

- Increase circulating power to 400kW circulating power
  - New laser amplifier (improve high-frequency sensitivity)
  - Point absorber free test masses (improve high-frequency sensitivity)
- Squeezing ~4.5 dB
  - Adaptive mode matching (improve broadband sensitivity)
  - Low-loss faraday isolator (improve broadband sensitivity)
  - Frequency dependent squeezing (FDS) (improve broadband sensitivity)
- Technical noise reduction
  - Stray light baffles (improve low frequency sensitivity)
  - Control systems improvements



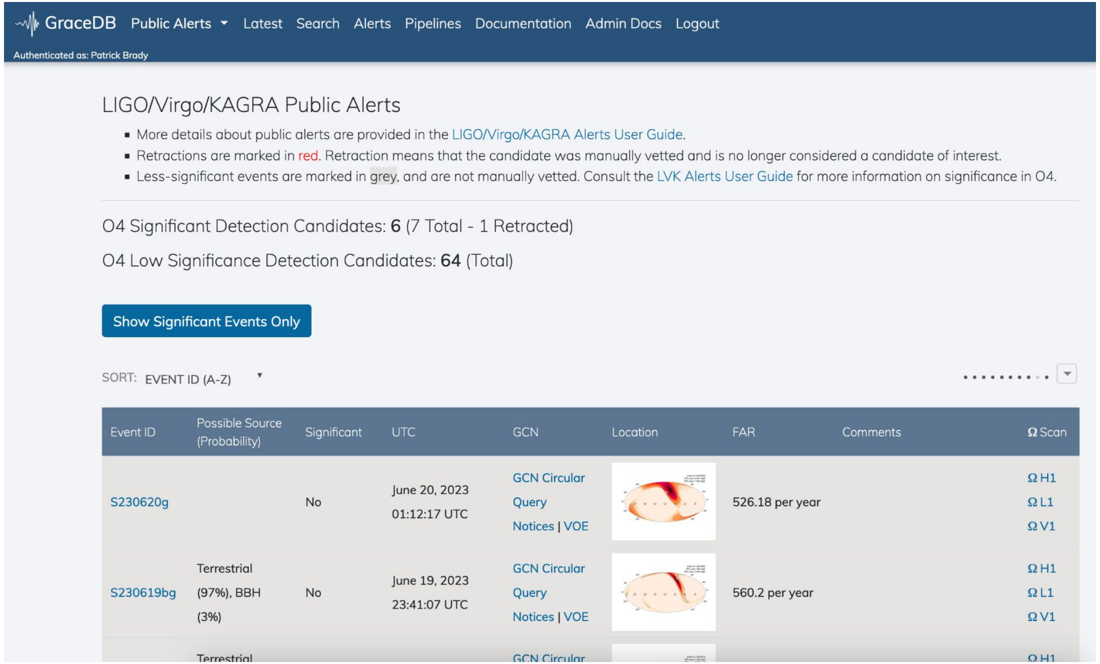




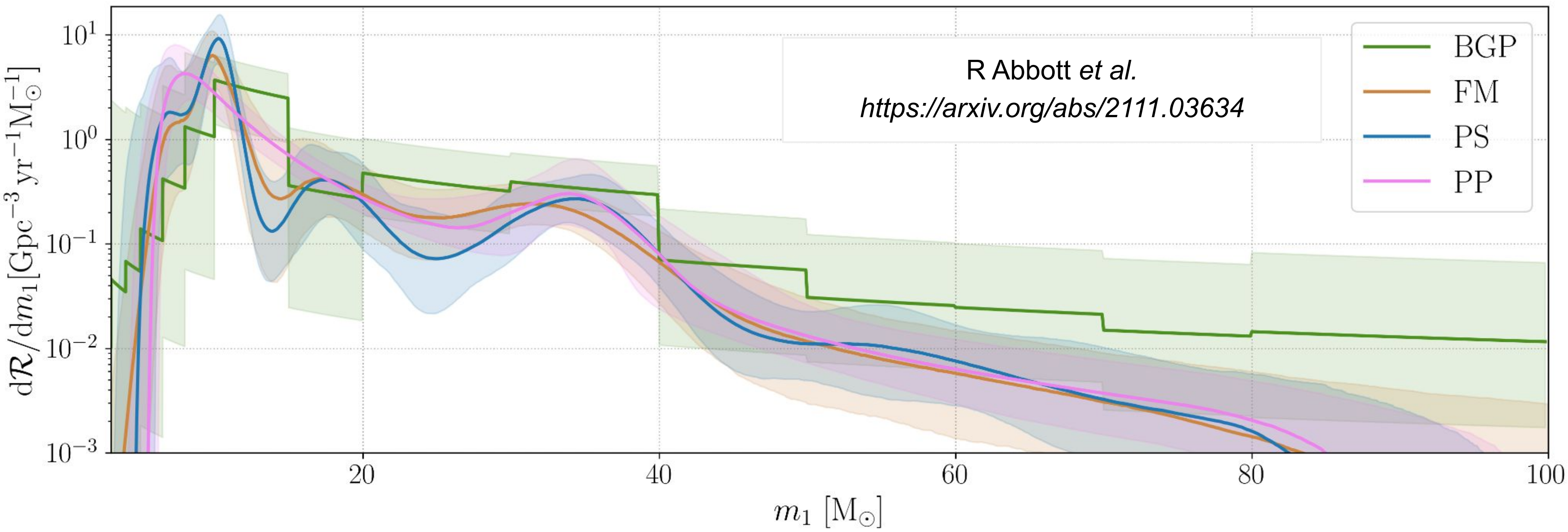


Back to observing!

- O4 started 24 May 2023: 20 months with up to 2 months commissioning
  - Virgo delayed due to damage to optics
- Binary detection rates
  - O3 ~ 1 / 5 days
  - O4 ~ 1 / (2-3 days)
- Improved public alerts
  - Localization
  - Classification
  - Latency
  - Early-warning alerts
  - Low-significance alerts



From one to many: measuring populations

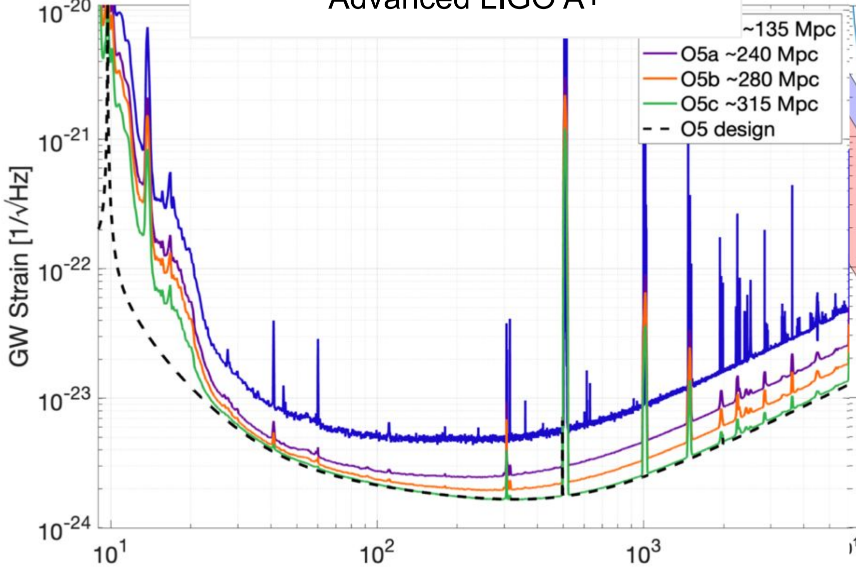


Merger rate density as a function of primary mass using 3 non-parametric models compared to the power-law+peak (pp) model.

Thanks to Dave Reitze & Giovanni Losurdo

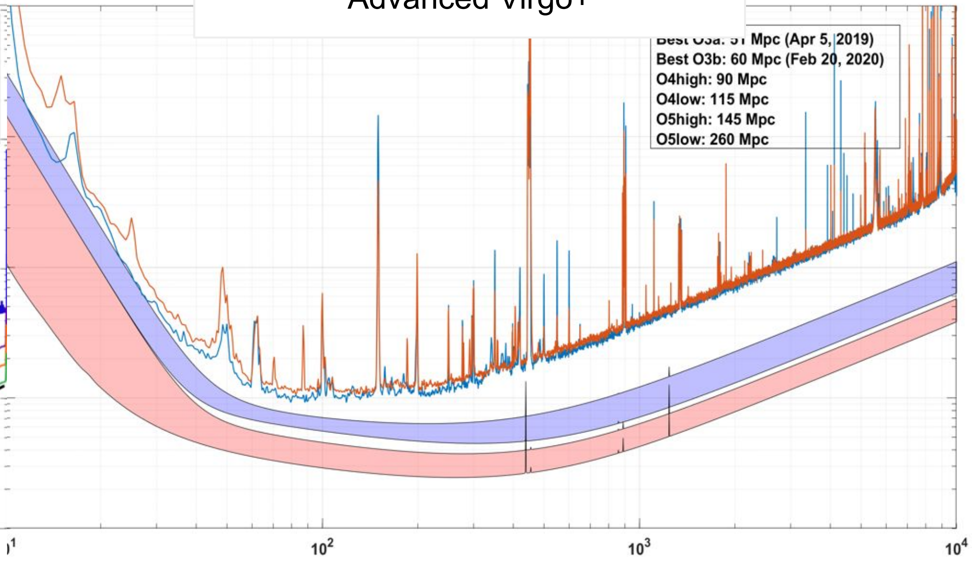
Working toward O5 sensitivity

Advanced LIGO A+



Frequency (Hz)

Advanced Virgo+



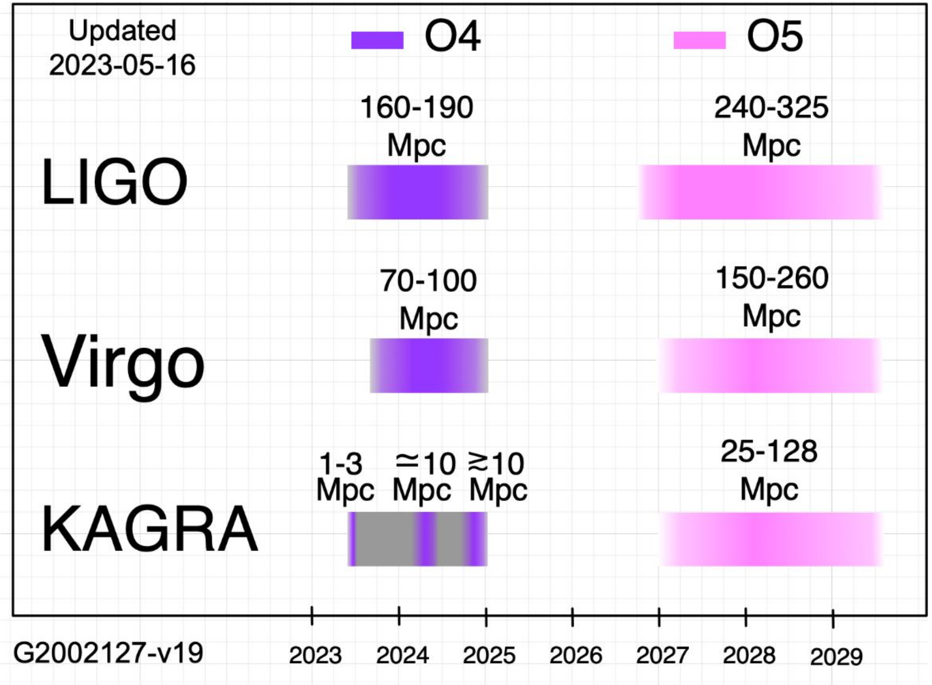
Frequency (Hz)

Full Power in the arm cavities: 750 kW
Frequency-dependent Squeezing* level of 6 dB
Test Masses with 2x lower coating thermal noise*

KAGRA will continue to work towards 130Mpc goal in O5

O5 Observing Run

- Current thinking
  - Start is paced by upgrades after O4: 1.5-2 years gap.
  - Intersperse commissioning and observations
- Binary detection rates
  - O3 ~ 1 / 5 days
  - O4 ~ 1 / (2-3) days
  - O5 ~ 3 / day
- Other science
  - Improved SNR
  - New sources?



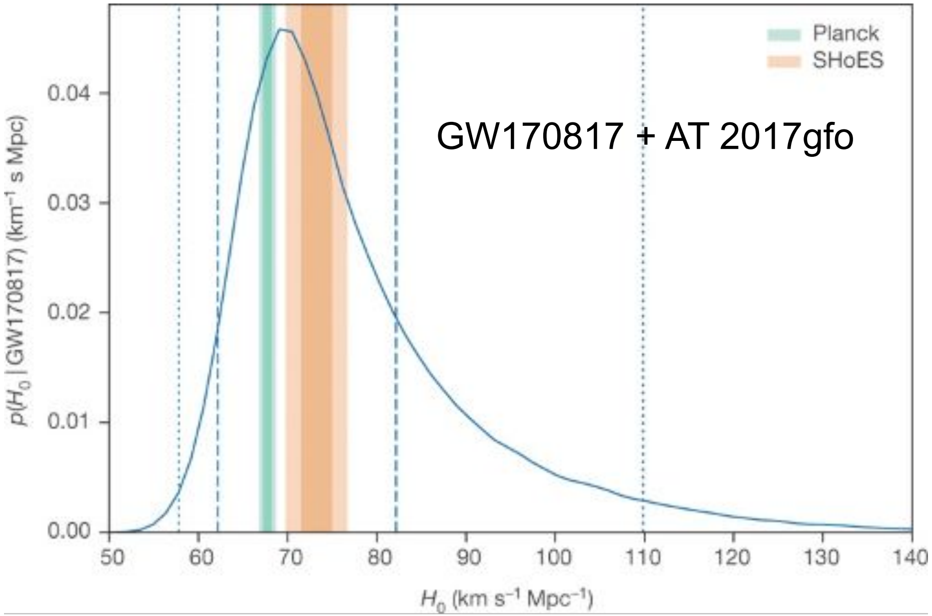
LIGO-Virgo-KAGRA anticipate observing to dovetail with next generation facilities

https://observing.docs.ligo.org/plan/

Cosmology with gravitational waves

- Gravitational waves from binaries are standard sirens
  - Measure the luminosity distance to the source and redshifted masses
  - Cannot measure redshift directly
- Get redshift some other way
  - Electromagnetic counterpart, e.g. GW 170817, GRB 170817A, AT 2017gfo
- Sub-percent accuracy with many
  - Cross correlate with galaxy redshifts [Schutz, Nature **323**, 310 (1986)]
  - Mass scale imprinted on spectrum of detected binary mergers [Will M. Farr et al 2019 ApJL 883 L42]

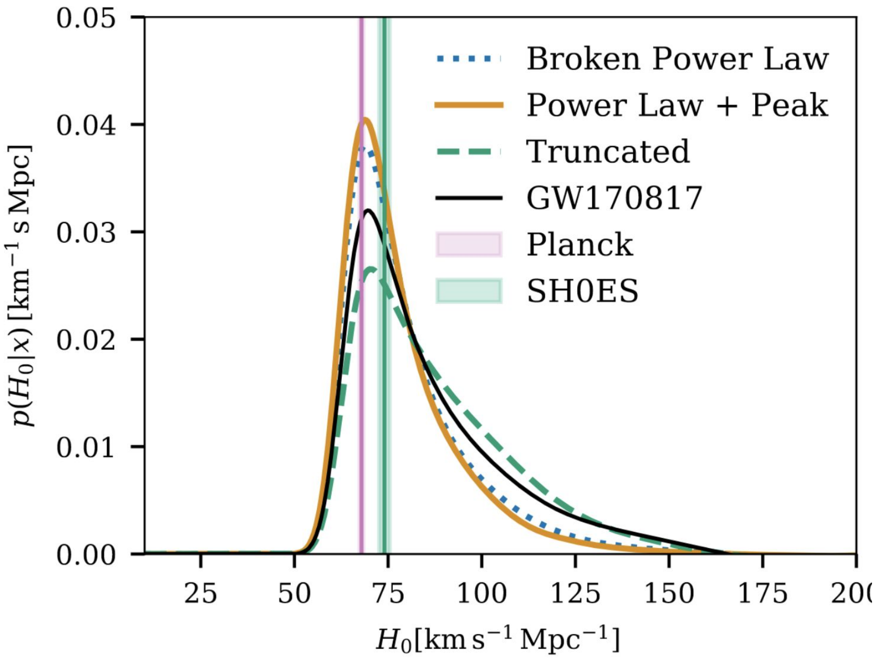
B P Abbott *et al. Nature* **551,** 85–88 (2017) doi:10.1038/nature24471



Challenges for cosmology with GW

- Binaries with detectable EM counterparts are rare
  - With ~5-10 BNS mergers detectable in O4, expect ~1 detectable kilonova.
  - GRBs further away, but only a fraction beamed to Earth.
- Sub-percent accuracy with many
  - Completeness of galaxy catalogs decreases rapidly with redshift.
  - Mass scales are highly uncertain, e.g. maximum black hole mass from PISN, or must be measured simultaneously.

R Abbott *et al. arXiv:2111.03604 (2021)*

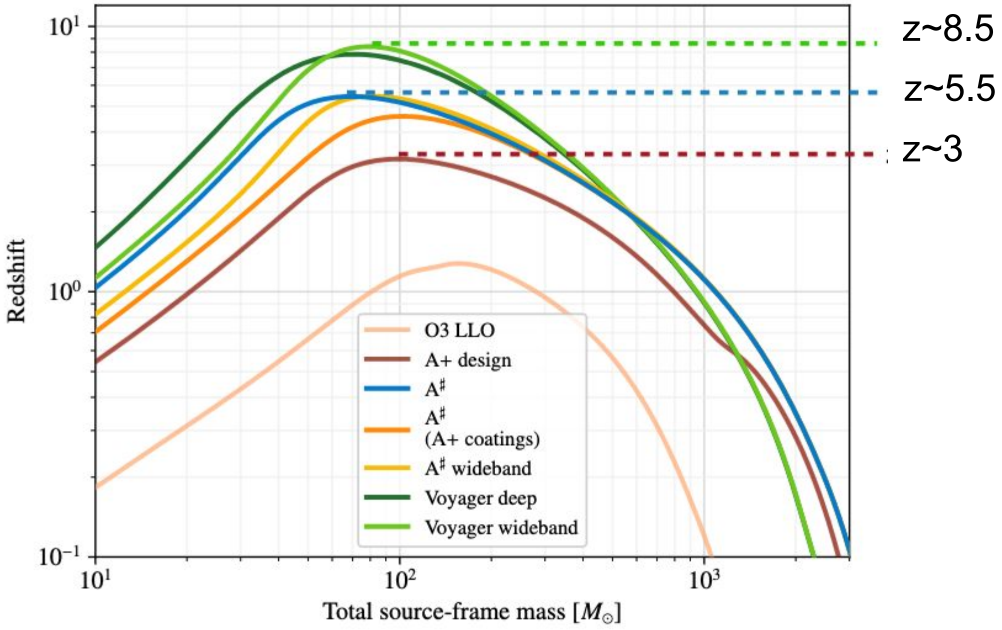


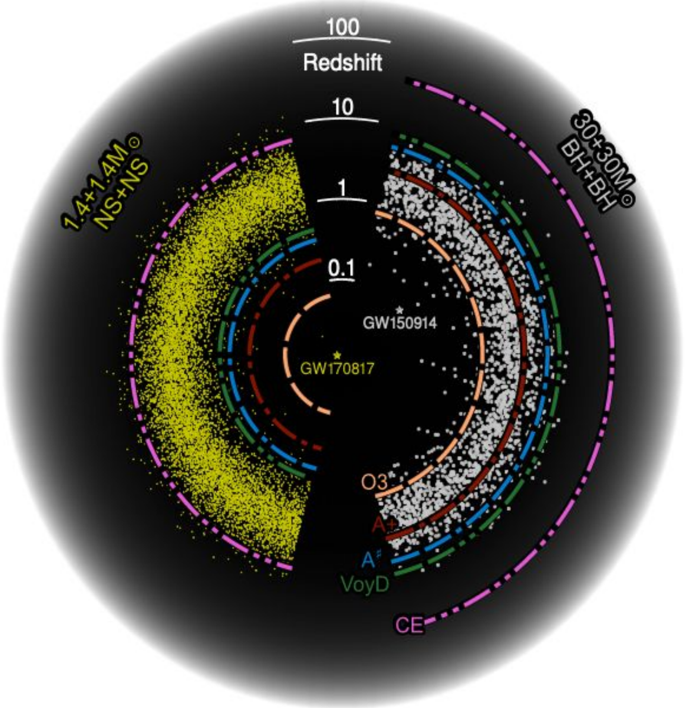
Early 2030s

- LIGO Aundha Observatory (LAO) is to be constructed in India and operated as part of the LIGO Observatories in the 2030s.
- $A^{\#}$: targeted improvements to the LIGO detectors
  - Report of LSC post-O5 study group [Fritschel et al, https://dcc.ligo.org/LIGO-T2200287/public]
  - Achieve close to a factor of 2 amplitude sensitivity improvement with larger test masses, better seismic isolation, improved mirror coatings, higher laser power, better squeezing …
  - Begin observing at the end of 2031 and observe for several years.
  - $A^{\#}$ an engine for observational science and a pathfinder for next-generation technologies.
  - A network including LIGO A# detectors would be a cornerstone for multimessenger discovery.
- Virgo has scoped similar improvements, called VirgoNEXT, with similar timetable. KAGRA is focused on reaching its current target.

Observational Science with A$^{\#}$

Horizon for optimally oriented and located binary mergers





See Fritschel et al, https://dcc.ligo.org/LIGO-T2200287/public

Observational Science with A$^{\#}$

- Probe the compact object binary population with unprecedented precision
  - Masses, spins, sub-populations.
  - Clues about their formation and astrophysical environment.

- Hubble constant measurement to sub-percent levels
- Black hole spectroscopy via sub-dominant modes
- Neutron star radius measurements to sub-km
- Enlarge discovery space: nearby supernova, continuous wave sources, stochastic background

Configuration	Annual Detections		
	BNS	NSBH	BBH
A+	135^{+172}_{-78}	24^{+34}_{-16}	740^{+940}_{-420}
A$^\sharp$	630^{+790}_{-350}	100^{+128}_{-58}	2100^{+2600}_{-1100}
A$^\sharp$ (A+ coatings)	260^{+320}_{-140}	45^{+60}_{-27}	1150^{+1450}_{-640}
A$^\sharp$ Wideband (A+ coatings)	200^{+250}_{-110}	40^{+54}_{-25}	970^{+1220}_{-540}
Voyager Deep	1280^{+1610}_{-710}	190^{+240}_{-110}	3100^{+3900}_{-1700}
Voyager Wideband	730^{+920}_{-410}	129^{+165}_{-74}	2300^{+2900}_{-1300}

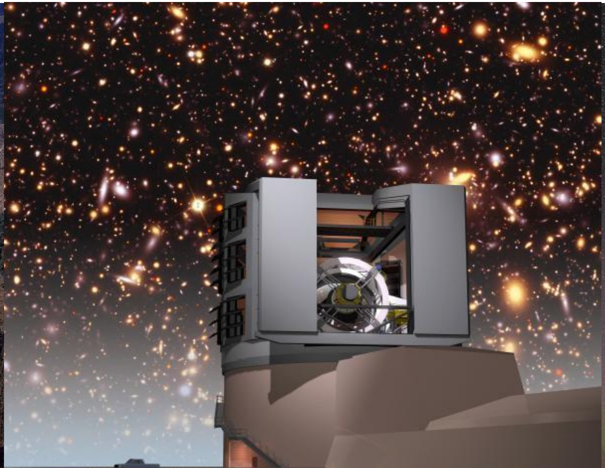
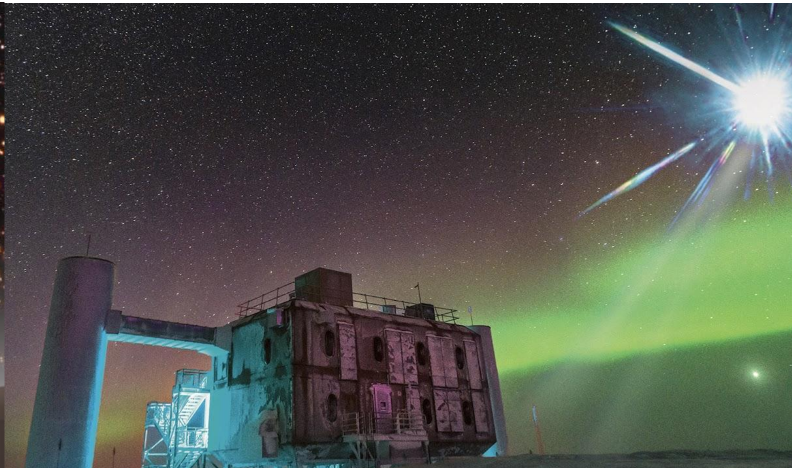
See Fritschel et al, https://dcc.ligo.org/LIGO-T2200287/public

LIGO is a cornerstone of MMA

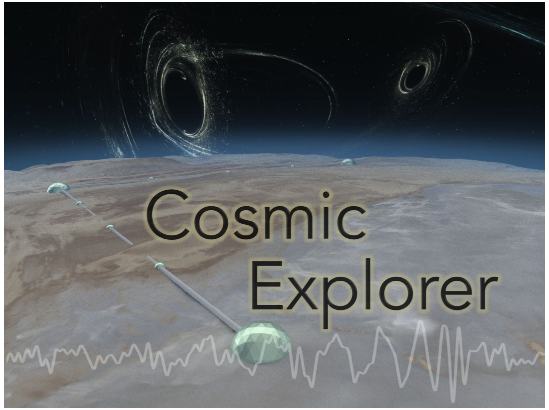
- The number of detections per year for four detector networks for binary neutron stars within z = 0.5

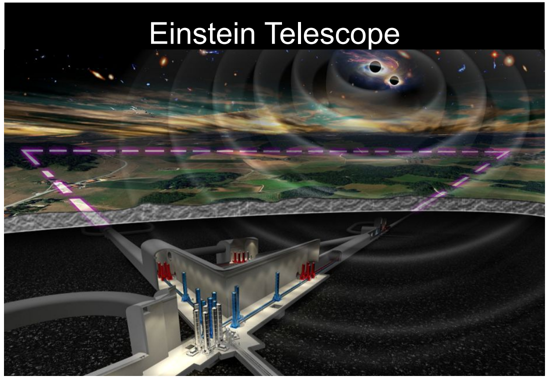
Metric	Ω_{90} (deg)^2		
Quality	≤ 100	≤ 10	≤ 1
$3A^\sharp$	$1.2^{+1.8}_{-0.9} \times 10^3$	$3.2^{+4.7}_{-2.5} \times 10^2$	$5.0^{+11.0}_{-5.0} \times 10^0$
$CE20 + 2A^\sharp$	$8.6^{+13.3}_{-6.4} \times 10^3$	$8.6^{+12.9}_{-6.8} \times 10^2$	$1.7^{+3.3}_{-1.5} \times 10^1$
$CE40 + 2A^\sharp$	$9.8^{+15.1}_{-7.3} \times 10^3$	$9.7^{+14.6}_{-7.6} \times 10^2$	$1.8^{+3.8}_{-1.6} \times 10^1$
$CE40 + CE20 + 1A^\sharp$	$1.4^{+2.1}_{-1.0} \times 10^4$	$3.4^{+5.3}_{-2.6} \times 10^3$	$9.7^{+15.7}_{-7.7} \times 10^1$

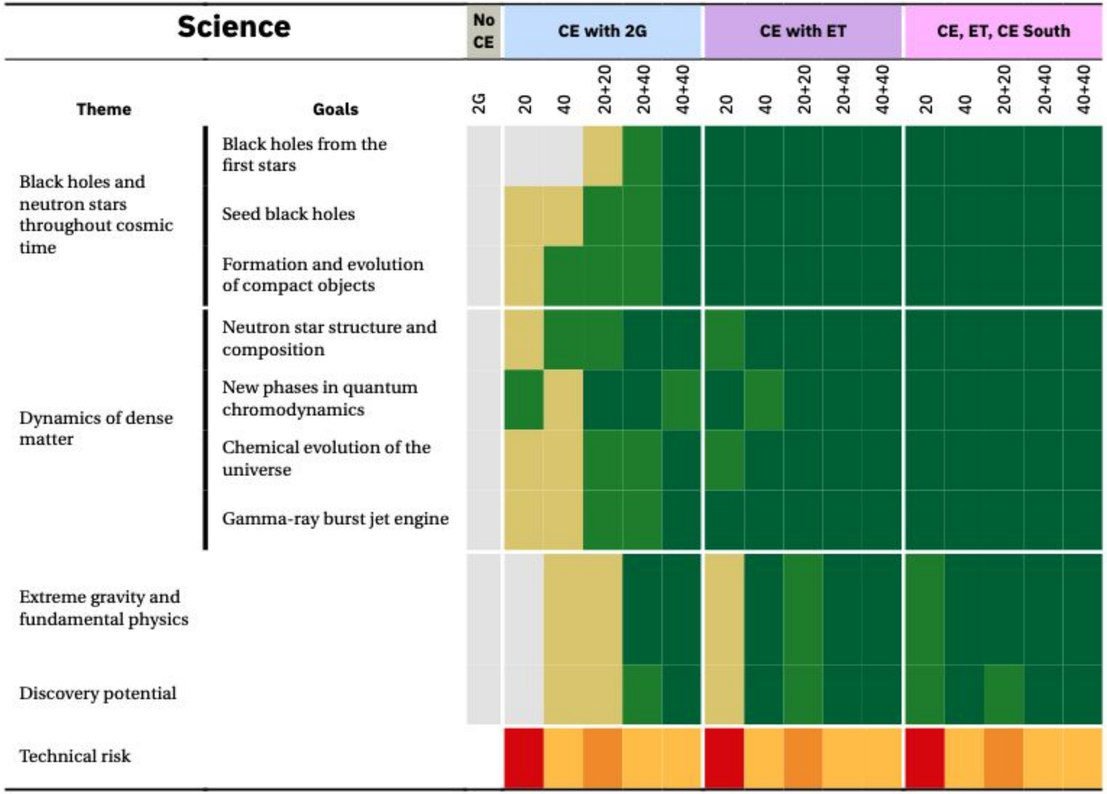
Ish Gupta & B. Sathyaprakash

Next Generation Detectors

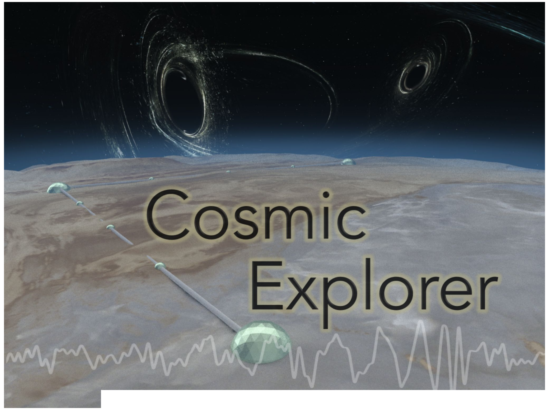






A Horizon Study for Cosmic Explorer
https://arxiv.org/abs/2109.09882

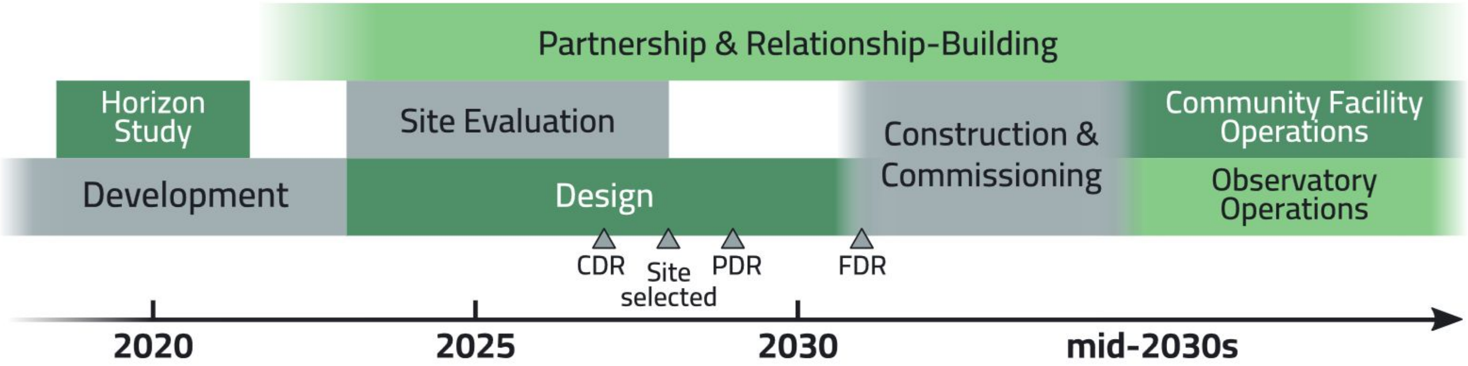
Cosmic Explorer Timeline



A Submission to the NSF MPSAC ngGW Subcommittee

https://dcc.cosmicexplorer.org/CE-P2300018/public

Top-level timeline showing a phased approach to design and construction.



Thank you!